

XRB continuum fitting with sensitive high-energy X-ray detectors

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ABSTRACT

The launch of the *Nuclear Spectroscopic Telescope Array* (*NuSTAR*) heralded a new era of sensitive high-energy X-ray spectroscopy for X-ray binaries (XRBs). In this paper we show how multiple physical parameters can be measured from the accretion disc spectrum when the high-energy side of the disc spectrum can be measured precisely using *NuSTAR*. This immediately makes two exciting developments possible. If the mass and distance of the source are known, the continuum fitting method can be used to calculate the spin and inner disc inclination independently of the iron line fitting method. If the mass and distance are unknown, the two methods can be combined to constrain these values to a narrow region of parameter space. In this paper we perform extensive simulations to establish the reliability of these techniques. We find that with high-quality spectra, spin and inclination can indeed be simultaneously measured using the disc spectrum. These measurements are much more precise at higher spin values, where the relativistic effects are stronger. The inclusion of a soft X-ray snapshot observation alongside the *NuSTAR* data significantly improves the reliability, particularly for lower temperature discs, as it gives a greatly improved measurement of the disc peak. High signal-to-noise data are not necessary for this, as measuring the peak temperature is relatively easy. We discuss the impact of systematic effects on this technique, and the implications of our results such as robust measurements of accretion disc warps and XRB mass surveys.

Key words: accretion, accretion discs – black hole physics – X-rays: binaries.

1 INTRODUCTION

There are two main methods for establishing the spin of a black hole X-ray binary (XRB). The first, the iron line or reflection method, relies on measuring the relativistic distortion of the narrow iron K α line induced by the rotation of the inner disc and by the gravitational redshift of the black hole (Fabian et al. 1989; Reynolds 2014). The second method, continuum fitting, relies on accurately establishing the luminosity of the disc spectrum, which is a strong function of the dimensionless spin parameter a (Zhang, Cui & Chen 1997; McClintock, Narayan & Steiner 2014). This method is generally only applicable when the mass, M , and distance, D , of the black hole concerned are known, as otherwise the luminosity and baseline temperature of the disc are unknown. The iron line method thus relies on precise measurement of the spectral shape and uses no

information about the absolute flux in the line, and the continuum fitting method largely relies on the flux and uses relatively little spectral information, due to the far smoother shape of the disc spectrum.

Both of these methods have been used extensively to measure the spin of X-ray binaries. Middleton (2016) updates the table of Miller & Miller (2015) and lists five XRBs with continuum fitting spins, eight with reflection spins, and six with both. Generally, the two techniques are applied to different spectra: continuum fitting requires a strong disc component, which can obscure or confuse the iron line measurements used in reflection spectroscopy. Because of this, the two methods are rarely used simultaneously, which limits how comparable their results are, as the accretion geometry may change between different observations. Despite this, the agreement between the two methods is generally good – only two sources have significant differences: 4U 1543–475 ($a_{\text{ref}} = 0.3 \pm 0.1$, $a_{\text{cont}} = 0.8 \pm 0.1$) and GRO J1655–40 ($a_{\text{ref}} = 0.98 \pm 0.01$, $a_{\text{cont}} = 0.7 \pm 0.1$). While it may be the case that this is just attributable to statistical or systematic errors, or truncation of the disc in some

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states, this may also be a signature of incorrect assumptions in the modelling, which can be used to probe accretion physics.

Because of the broad spectral shape of the disc it is much harder to measure the relativistic distortion of the spectrum, and instead spin measurements using continuum fitting rely on measuring the temperature and luminosity of the disc (see reviews by Remillard & McClintock 2006; McClintock et al. 2014). Both of these parameters are strongly affected by the spin, which controls the radius of the innermost stable circular orbit (ISCO). The bulk of the high-energy emission from the disc comes from the innermost radii, where the disc is hottest and the radiation can be strongly relativistically boosted. However, this also depends on the inclination of the disc with respect to the observer, as this controls the level of boosting and Doppler shifting of the emitted radiation. Again, this is most prominent at small radii, where the orbital velocities are largest. This degeneracy is broken by assuming that the inclination of the inner disc is the same as that of the binary system, which can be established along with M and D by using dynamical measurements (see e.g. Özel et al. 2010, and references therein). This technique has successfully been applied to a large number of spectra from several different sources (e.g. McClintock & Remillard 2006; Shafee et al. 2006; Liu et al. 2008; Gou et al. 2009, 2011; Steiner et al. 2009b, 2010, 2011).

However, the requirement for a known mass, distance, and inclination limits the application of the continuum fitting method to the handful of sources where these parameters are known, and relies on the assumption that the disc and binary inclinations are the same. In principle, however, this limitation can be partly overcome with sufficiently high-quality data. While there is a strong degeneracy between a and i , it is not absolute. This means that with enough signal a and i can be measured simultaneously. This necessitates the use of precision spectroscopy with a highly sensitive instrument, which covers the high-energy side of the disc spectrum.

NuSTAR (Harrison et al. 2013) is the first and currently only focussing hard X-ray telescope (HXT), although it was briefly joined by the HXT aboard *Hitomi* (Kokubun et al. 2012). The *NuSTAR* detectors are uniquely suited to observations of X-ray binaries due to the absence of the pileup effects that plague CCD detectors at high count rates (when multiple photons arrive within a single read-out time and are interpreted as one event, see e.g. Miller et al. 2010). By reading each event as it occurs, rather than reading all events with a fixed frequency, this problem is eliminated. This means that highly sensitive measurements of the broad iron line have been made possible by *NuSTAR* (Miller et al. 2013a,b; King et al. 2014; Tomsick et al. 2014; Fürst et al. 2015; Parker et al. 2015, 2016; Walton et al. 2016, 2017).

This paper is a follow-up to Parker et al. (2016), hereafter P16, where we demonstrated the potential of *NuSTAR* to constrain multiple disc parameters simultaneously. In this work, we will examine the suitability of continuum fitting to constrain both spin and (inner disc) inclination at the same time, using simulations. In Section 2 we describe the basic principle of fitting the disc spectrum for multiple parameters. In Section 3 we simulate a large number of spectra to test the reliability of the method. In Section 4 we explore some potential sources of systematic error in this technique. Our discussion and conclusions are in Sections 5 and 6.

2 PRINCIPLE OF THE METHOD

In Fig. 1 we show the data/model residuals of the very high state spectrum of GX 339-4 to the best-fitting model presented in Parker et al. (2016), with the relativistic disc spectrum removed. It is clear

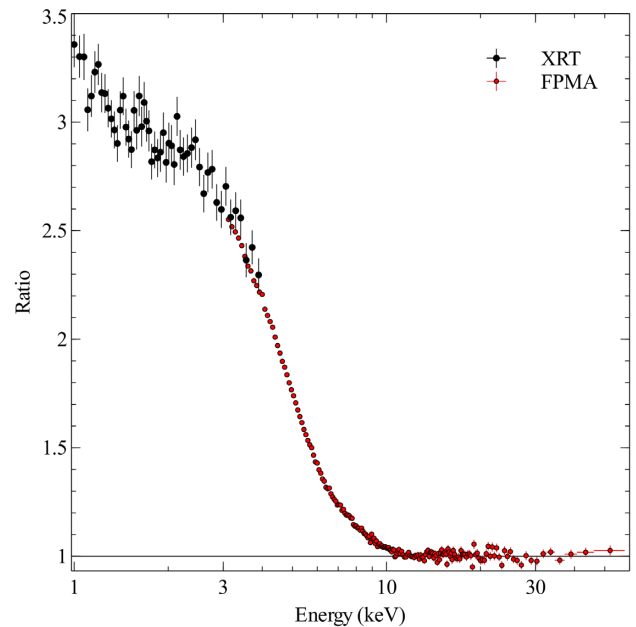


Figure 1. The data/model ratio from the GX 339-4 very high state *NuSTAR* and *Swift* observation P16, with the *kerrbb* disc model removed. These residuals show the disc contributes a significant amount of flux up to 10 keV. Note that we only show the spectrum from one of the two *NuSTAR* modules, so twice as much data are available. The peak of the disc spectrum is around 1 keV, but the disc fraction continues to rise below this point due to the drop in photons in the Compton-scattered tail.

from this figure that the high-energy side of the disc is contributing a significant amount of flux well into the *NuSTAR* band (up to ~ 10 keV), meaning that small distortions in the spectral shape can potentially be measured, given the high quality of the *NuSTAR* data (only one of the two FPM spectra is shown, and below ~ 15 keV the error bars are smaller than the points).

To explore how this high sensitivity can be exploited, we consider a relativistic disc spectrum, using the KERRBB model (Li et al. 2005). We simulate an idealized 0.5–10 keV spectrum, with no background, no noise, a 100 ks exposure time, and the *XMM-Newton* instrument response, using KERRBB with parameters $M = 10 M_{\odot}$, $D = 10$ kpc, $a = 0.998$, $i = 40^{\circ}$, and $\dot{M} = 0.2 \dot{M}_{\text{Edd}}$, where \dot{M}_{Edd} is the Eddington accretion rate. We then re-fit this data, stepping a from 0.998 to 0, with M and D fixed and the other parameters free to vary. In this case, as in all others, we keep the normalization of KERRBB fixed to 1. The results of this are shown in Fig. 2. It is obvious that at low energies the spectra are extremely similar. However, at higher energies there are very large differences between the models (50 per cent less flux at 10 keV in the $a = 0$ case). These differences are largely due to the increased high-energy flux for rapidly rotating black holes, caused by Doppler shifting and relativistic beaming of the photons coming from the approaching side of the disc. At higher spin values, the velocities at the ISCO are larger, meaning that the beaming is stronger. This can be mimicked to an extent by increasing the viewing inclination (with respect to the disc axis, so that the disc is more edge-on to the observer) giving a higher line-of sight velocity. However, this does not affect the absolute velocity of the disc, and so cannot fully reproduce a higher spin spectrum.

In relative terms, the differences in spectral shape measured using this method are small, due to the rapid drop in disc flux at high energies, and can only be reliably measured because of the very

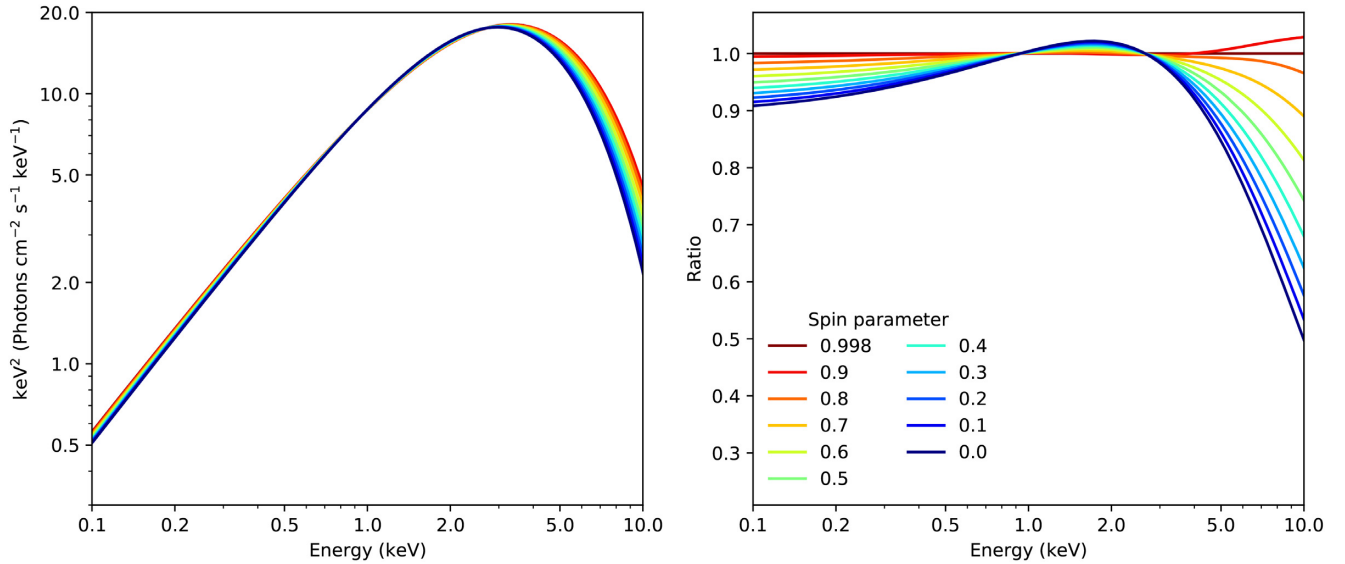


Figure 2. Left: KERRBB models with fixed a fit to a simulated spectrum with $a = 0.998$. Below ~ 3 keV the spectra are almost completely degenerate, but at high energies large differences start to appear. Right: As left, but as a ratio to the original $a = 0.998$ model. At 10 keV, 50 per cent of the model flux is lost by fitting with $a = 0$. This can easily be measured, provided that the disc spectrum can still be constrained at high enough energies.

Table 1. Estimated parameters and fit statistics for fitting a spectrum with the spin fixed at different values.

Parameter	a	i (deg)	\dot{M} (10^{18} g s^{-1})	χ^2/dof
Simulated	0.998	40	2.74	
Best-fit	>0.98	43^{+2}_{-3}	$3.2^{+0.2}_{-0.4}$	383/371
Fixed	0.9	54.70 ± 0.03	4.770 ± 0.001	418/372
	0.8	63.52 ± 0.03	6.711 ± 0.002	465/372
	0.7	70.36 ± 0.03	9.400 ± 0.005	504/372
	0.6	75.70 ± 0.03	13.21 ± 0.01	517/372
	0.5	79.74 ± 0.03	18.53 ± 0.02	519/372
	0.4	82.71 ± 0.02	25.71 ± 0.04	526/372
	0.3	84.88 ± 0.02	35.06 ± 0.05	567/372

high sensitivity of *NuSTAR* or similar instruments. To demonstrate this, we consider a more realistic simulated *NuSTAR* spectrum from our sample described in Section 3.1, with a power-law tail, where the spin is well constrained. We take the spectrum for $M = 10 M_{\odot}$, $D = 3$ kpc, $a = 0.998$, and $i = 40^\circ$, and manually fit it. We then re-fit with the spin fixed at different values to determine the effect on the spectrum. The results of this are given in Table 1. While this is a simplified spectrum, the result is clear: while there is a degeneracy between spin and inclination, a high-spin relativistic disc spectrum cannot be approximated by a low-spin disc spectrum to arbitrary precision, and the level of spectral precision required to break this degeneracy is well within the reach of current instrumentation. There is therefore no reason, in principle, why spin and inclination cannot be simultaneously measured.

The most obvious application of this is to free the inclination parameter, allowing it to differ from that of the binary system. Of the three fixed parameters, this is the only one that could theoretically differ from that obtained from dynamic measurements. When the black hole spin axis and the binary orbit are not aligned, the inner disc is likely to warp (Bardeen & Petterson 1975), aligning itself with the black hole spin, while the outer disc remains aligned with the binary system. There have been hints of differences in

inclination between the dynamic and reflection measurements in some sources. For example, inclinations measured using reflection spectroscopy of Cyg X-1 with *NuSTAR* (Tomsick et al. 2014; Parker et al. 2015; Walton et al. 2016) have generally been $10\text{--}15^\circ$ higher than that inferred of the binary system (Orosz et al. 2011). Using this technique we can now obtain an independent measurement of the inclination using continuum fitting, giving three independent measurements of the inclination (binary, reflection, and continuum). If those obtained from the X-ray spectra (which only measure the inner disc) systematically differ from the binary inclination (which determines the outer disc inclination), we can be much more confident of a warp in the disc.

A second, less obvious, application is to combine the continuum fitting and reflection methods to swap which parameters are fixed and which are measured. Instead of using M and D as the input for the continuum fitting model, the a and i values obtained with reflection can be used, and M and D can be constrained. This method has two main sources of error: the intrinsic uncertainty in measurements of M and D , due to the degeneracy between the two and the accretion rate, and the uncertainty inherited from the a and i values from reflection. How reliably this method can be used depends on the relative strengths of these effects – ideally, the intrinsic uncertainty will be small, so the net error is dominated by the uncertainty in a and i from reflection, which can be measured relatively accurately. We explore this in Section 3.3.

3 SIMULATIONS

3.1 Spin and inclination

To test the accuracy of fitting spin and inclination simultaneously we simulate a grid of spectra with different a , i , M , and D values and then fit the model back to the data. For this test, we use a simple model of a relativistic disc (POLYKERRBB, a version of KERRBB modified to allow high-order interpolation, see Appendix A) and Comptonized tail (SIMPL), modified by absorption (TBABS) with a column density of $5 \times 10^{21} \text{ cm}^{-2}$. We simulate 41 600 spectra,

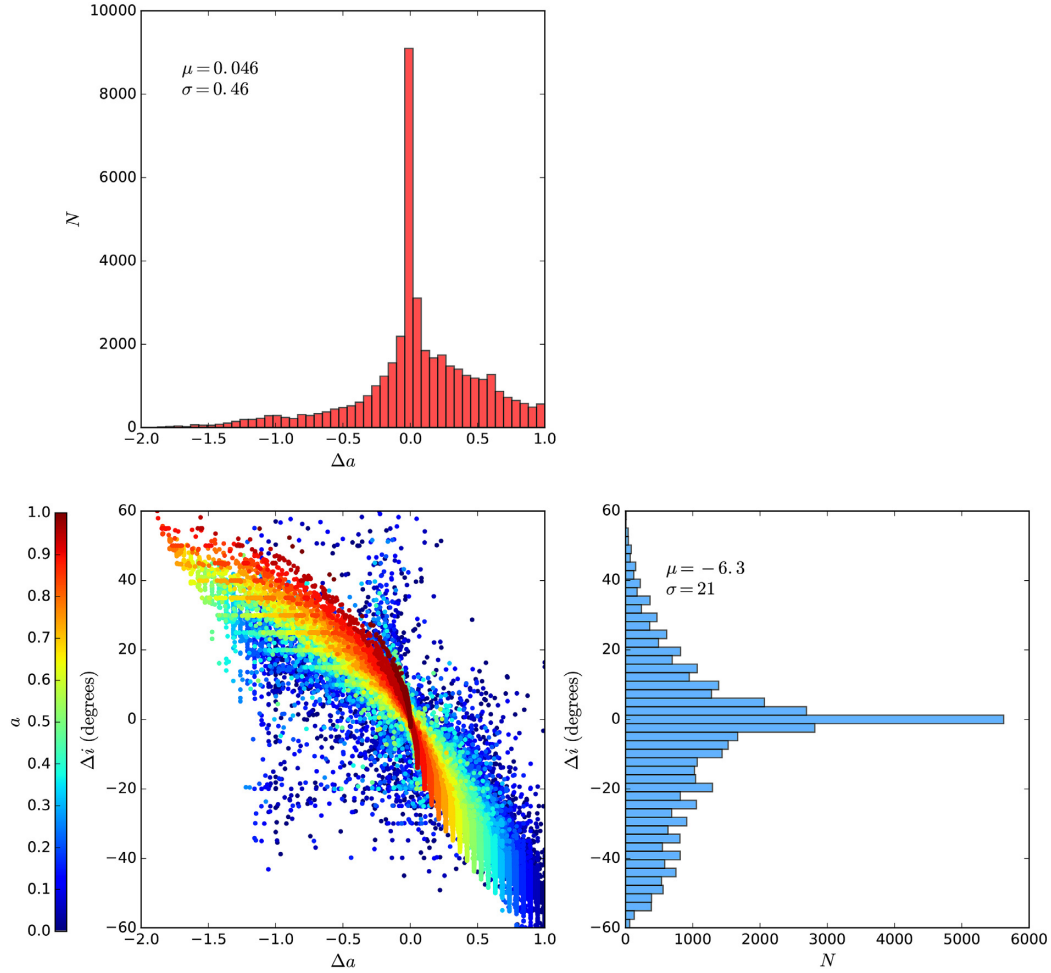


Figure 3. The spin-inclination degeneracy for all 41 600 spectra in the sample. The lower left plot shows the difference between the measured and simulated values of a and i , colour-coded by the spin, and the two histograms show the distributions in Δa and Δi after marginalizing over the other parameter. μ and σ are the mean and standard deviation, for each parameter. The series of peaks in the inclination histogram are caused by fits hitting the edge of the allowed range of inclinations.

with 16 values of M from 5 to 20 M_{\odot} , 10 values of D from 1 to 10 kpc, 20 values of a from 0 to 0.998, and 13 values of i from 20° to 80° . All grid parameters are linearly spaced. The photon index of SIMPL is set to 2, and the fraction of photons scattered into the power-law tail is 0.1. At each step on the grid, the accretion rate is adjusted to $0.2 \times \dot{M}_{\text{Edd}}$. We simulate a long 100 ks *NuSTAR* exposure¹ for each spectrum, using the official background and response files for simulations.² We then fit these spectra with the same model, fixing the mass and distance at each step but allowing all other parameters (including those of SIMPL) to vary freely. We then record the final best-fitting spin and inclination estimates, and their errors. Additionally, we record the 3–10 keV flux of the disc component so that we can control for the effect of flux changes caused by the various parameters.

Due to the large number of spectra we use, the procedure for simulating and fitting had to be automated. We use the PYXSPEC interface for XSPEC, which allows it to be called from PYTHON.

¹Note that for our purposes, exposure time is entirely degenerate with distance, and thus the effect of shorter exposures will be identical to the effect of larger distances.

²https://www.nustar.caltech.edu/page/response_files

We use two scripts, one which simulates the spectra based on the specified parameters, and one which fits the resultant fake data. We manually verified a small number of randomly selected fits to ensure that the program was functioning as expected. Each simulation or fit was run individually, distributed over the X-ray cluster in Cambridge using HTCONDOR. For the sake of speed, we simulate only one FPM spectrum, but double the exposure time to compensate. We rebin each spectrum to have a signal-to-noise ratio of 50 in every bin, and to oversample the data by a factor of 3. We ensure that for both simulating and fitting the energy grid is extended sufficiently far beyond the boundaries of the *NuSTAR* response that all models function as intended.

When fitting, we first set the parameters to those of the simulation, then run the standard XSPEC fit command. After the fit completes, we calculate the errors on a and i (frequently finding a new best fit in the process) using the standard XSPEC algorithm.

Fig. 3 shows the difference between the simulated and measured values of a and i for all 41 600 spectra (Δa and Δi , respectively). The distributions of a and i are both strongly peaked at the correct value ($\Delta a = \Delta i = 0$). For a given spin value, indicated by a single colour, the points are tightly confined to a curve showing the degeneracy between a and i , with some slight differences caused by different M value. The lack of scatter about the curves indicates how strong a

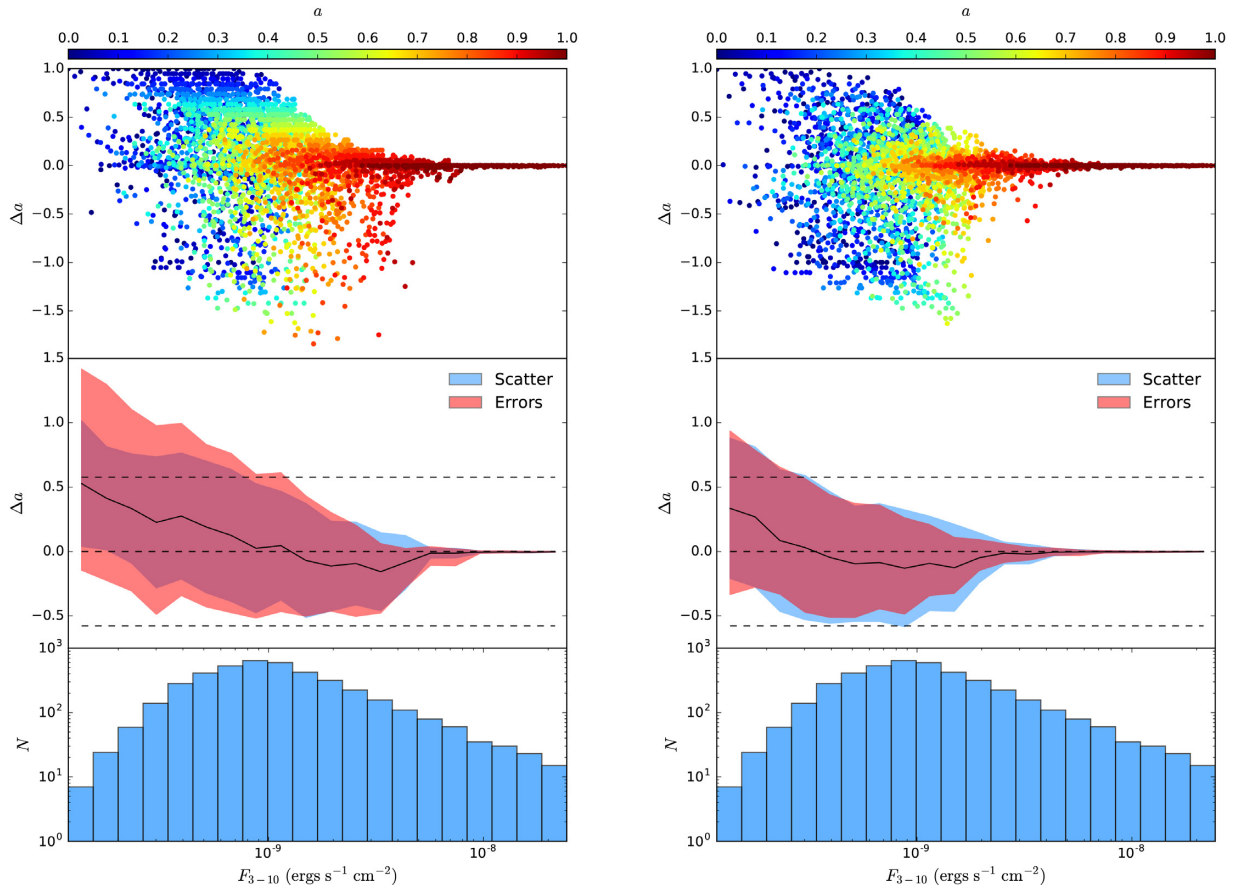


Figure 4. *Left:* The accuracy of the spin measurements as a function of flux. Top panel: All 4160 points for $D = 5$ kpc, colour-coded by spin. Middle: The scatter in those points (blue) and the mean error for each point (red). Dashed horizontal lines indicate the standard deviation expected from a uniform distribution between ± 1 (i.e. if the spin was unconstrained and selected at random). Bottom: Histogram of the number of points in each flux bin. *Right:* As left, but for joint *NuSTAR* XRT fits.

constraint can be obtained by *NuSTAR* alone on a single parameter - if the inclination is fixed, then the spin can be measured very precisely. It is obvious from this figure that the constraints are stronger at high spin, with the points spread around the degeneracy curve at low spin and tightly clustered at maximal spin. This is caused by three separate effects: the disc extends further into the *NuSTAR* bandpass at high spin; the changes in spectral shape are much more pronounced; and the signal to noise is higher due to high spin spectra being brighter.

In the left-hand panel of Fig. 4 we show the difference between the simulated and measured a values, plotted as a function of 3–10 keV flux, for all spectra at $D = 5$ kpc. There is a clear trend with flux: the higher the flux, the better the constraint obtained.³ Below around 10^{-9} erg s $^{-1}$ cm $^{-2}$, the scatter is around the magnitude expected for randomly chosen values, and there is a systematic bias towards higher spin. However, this increased uncertainty at low fluxes is reflected in the increased size of the errors (shown in the middle panel of Fig. 4). Based on this, we conclude that the measured errors on a (and hence i , as the two are strongly correlated) are likely to be accurate, so the true spin value is likely to be within the measured error interval.

³Note that, in the case where the background is low, the source flux is interchangeable with exposure time or the effective area of the instrument used.

We also investigate the uncertainty in the measured spin as a function of spin (Fig. 5, left). The top panel is a 2D histogram, showing the distribution of measured spins as a function of spin. For all values of spin, there is a peak at the correct value, but there are also many points clustered at the parameter limits. Generally, the higher the spin, the more accurate the measurement of spin. Similarly to Fig. 4, the measurement errors accurately reflect the scatter in the values, meaning that measured values are likely to be reliable.

A secondary effect of increased spin, which is particularly important for *NuSTAR* data, is that it shifts the peak of the disc spectrum to higher energies, bringing it into the *NuSTAR* band. In Fig. 6, we show the power-law residuals of four simulated spectra with different spin values but otherwise equivalent parameters. These flux changes and energy shifts are clearly visible – the highest spin spectra are both much brighter and extend much further into the *NuSTAR* band, making it far easier to constrain their spins.

3.2 The effect of soft X-ray data

To investigate the effect of including limited soft X-ray data, we simulate a corresponding 5 ks Neil Gehrels *Swift* Observatory X-ray Telescope (*Swift* XRT) snapshot spectrum for each *NuSTAR* spectrum. The simulated *Swift* XRT spectra are rebinned to a signal-to-noise ratio of 30. We then fit using the same procedure as before.

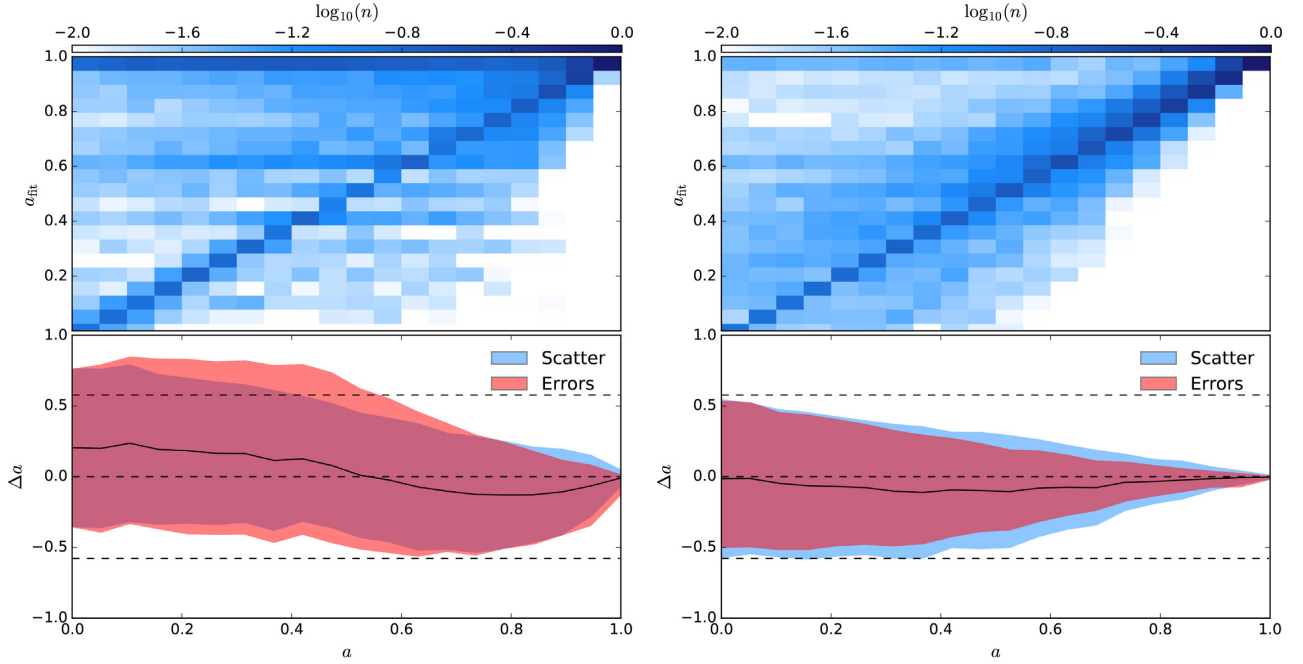


Figure 5. *Left:* The accuracy of the spin measurements as a function of spin. Top panel: 2D histogram, showing the spin measurements for all 41 600 spectra as a function of spin. There is a strong peak at the correct value for all spins, but a significant number hit the upper limit for all values of spin. Bottom panel: The standard deviation of the measured spins as a function of simulated spin, and the mean 1σ error (calculated with the XSPEC error command) as a function of spin. The two are a good match, implying that the XSPEC errors are reliable. *Right:* As left, but for the case where a *Swift* XRT snapshot is used in addition to the *NuSTAR* spectrum. The correlation is tighter, less biased, and has smaller scatter. Most notably, far fewer measurements hit the upper limit.

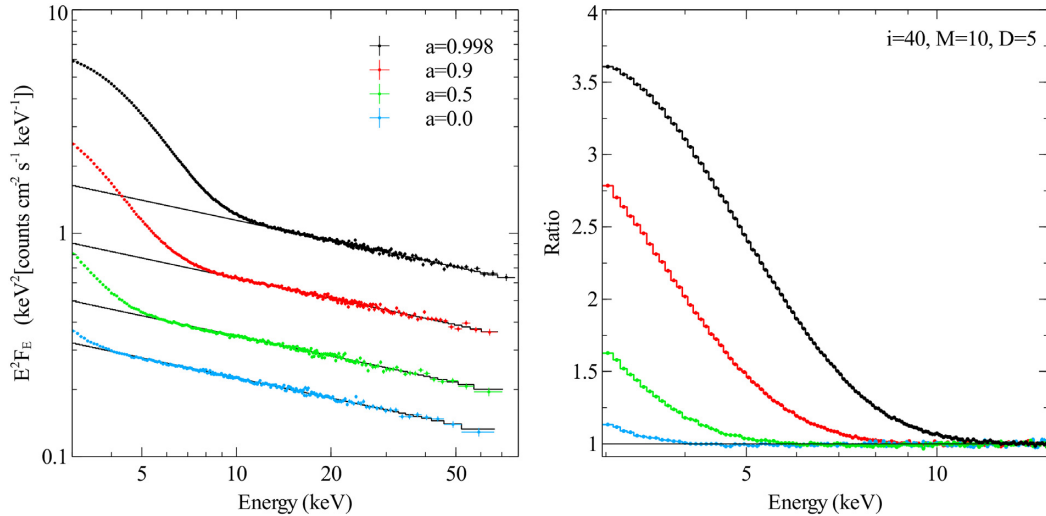


Figure 6. *Left:* Unfolded disc plus Comptonization spectra for a range of spin values at the same accretion rate, taken from our simulated spectra. Data are fitted with a power law in the 20–79 keV band, which is then extrapolated to lower energies. *Right:* Ratio of the data to the extrapolated power law below 20 keV.

The inclusion of XRT data, while it is relatively poor quality compared to the *NuSTAR* data, makes a significant improvement to the results. In Fig. 7, we show the spin/inclination degeneracy for these spectra, which has a much smaller scatter than that in Fig. 3 for the *NuSTAR* data alone. Importantly, the inclusion of XRT data also removes the bias in inclination, reducing the mean Δi from -6.3° to -0.26° , well below the measurement error.

We also show the effect of including XRT data on the spin scatter as a function of flux and spin (right-hand panels of Figs 4 and 5, respectively). In both cases, the inclusion of a short XRT exposure

greatly improves the results, extending the validity of the method to lower fluxes and lower spin values.

While these additional spectra contribute very little to the total signal, they allow the peak of the disc spectrum to be accurately measured, which is crucial when the disc spectrum does not extend far enough into the *NuSTAR* band for this curvature to be seen. This means that the effective temperature can be determined from the *Swift* snapshot, while the flux and the relativistic distortion can be measured by *NuSTAR*, giving a significantly more accurate result.

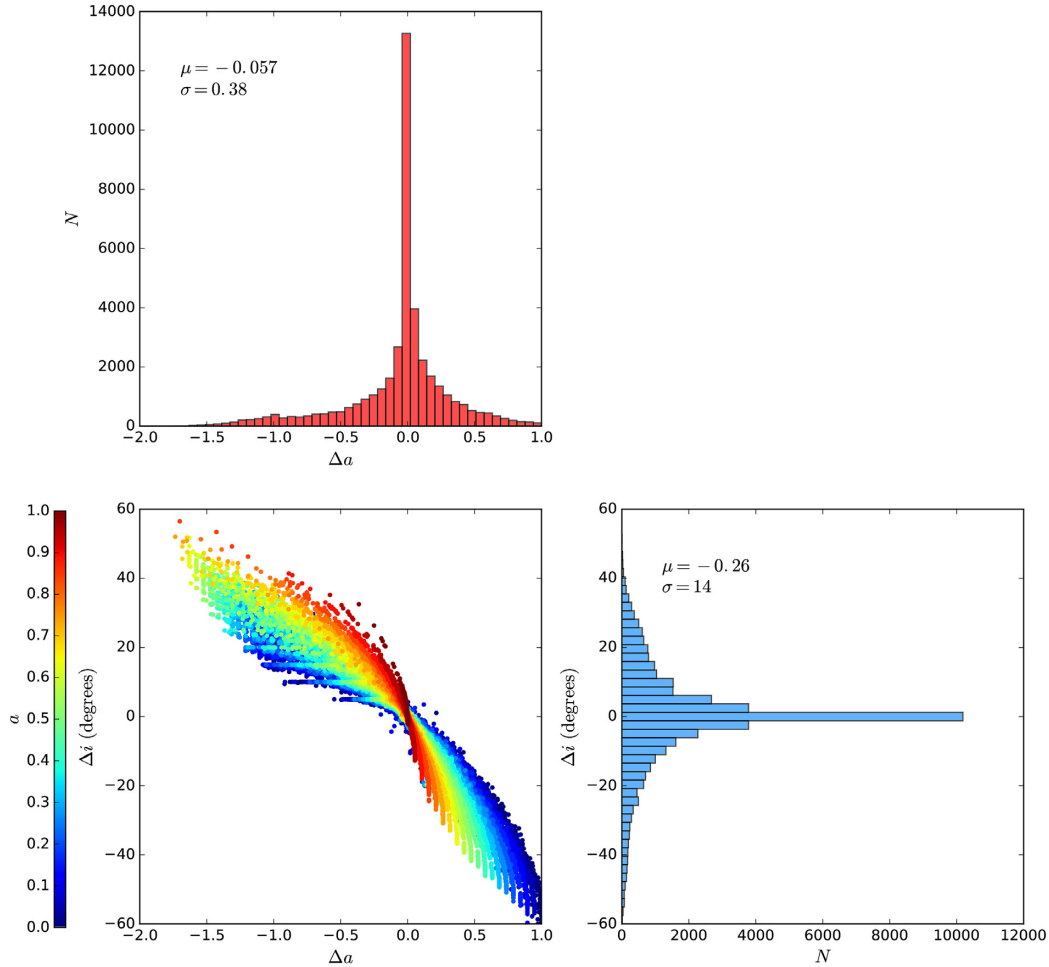


Figure 7. The spin-inclination degeneracy for all 41 600 spectra in the sample, as in Fig. 3, but with a 5 ks *Swift* XRT snapshot included along with each *NuSTAR* exposure. The lower left plot shows the difference between the measured and simulated values of a and i , colour-coded by the spin, and the two histograms show the distributions in Δa and Δi after marginalizing over the other parameter. As far fewer fits hit the edge of the parameter space, the peaks visible in the inclination histogram in Fig. 3 are not visible here. Both distributions are more strongly peaked at the correct value ($\Delta a = \Delta i = 0$), and the scatter in the points is much smaller.

3.3 Mass and distance

Having established that the continuum method can constrain two physical parameters of interest, we now move on to investigating M and D constraints. As in the case of spin and inclination, there is a strong degeneracy between mass and distance (and accretion rate). In this case, however, there is no extra spectral information to break this degeneracy, and this set of parameters is perfectly degenerate by definition within the model. This means that, contrary to the result of P16, M , D , and \dot{M} cannot be simultaneously constrained (see Section 5 for discussion of this).

To demonstrate this, we consider an individual simulated spectrum ($M = 10 M_{\odot}$, $D = 5$ kpc, $a = 0.9$, $i = 60^{\circ}$). We fix a and i at the simulated values, then step M from 1 to 100 M_{\odot} in 1000 steps, fitting for D at each mass. The results of this are shown in Fig. 8. There is a perfect degeneracy between M and D – the χ^2 value is constant across the whole range of values, so there is no way to distinguish between any two points statistically. This degeneracy is extremely well constrained – the chi-squared increases massively even a small distance from the relation (error bars are shown on the plot, and are much smaller than the points), so the ratio between the parameters is fixed and easily measurable. This means that if any

one of the parameters M , D , or \dot{M} can be constrained by another method, then the other two are immediately constrained to the same degree of accuracy. There are two main scenarios where one of these parameters could be established independently: measuring the distance to the companion, for example using the *Gaia* parallax, or estimating the accretion rate from the hardness–intensity diagram.

4 SOURCES OF SYSTEMATIC ERROR

Our results so far have shown that the a and i can be independently measured, and that the ratio between M and D can be easily and strongly constrained. In both cases, the statistical errors are an accurate reflection of the degree of uncertainty in our simulated data. However, this does not address the likely systematic error, which may be large and even dominate in some circumstances. In this section we attempt to establish some of the largest potential sources of systematic error, where they are particularly applicable to these techniques or have not been extensively discussed by previous authors.

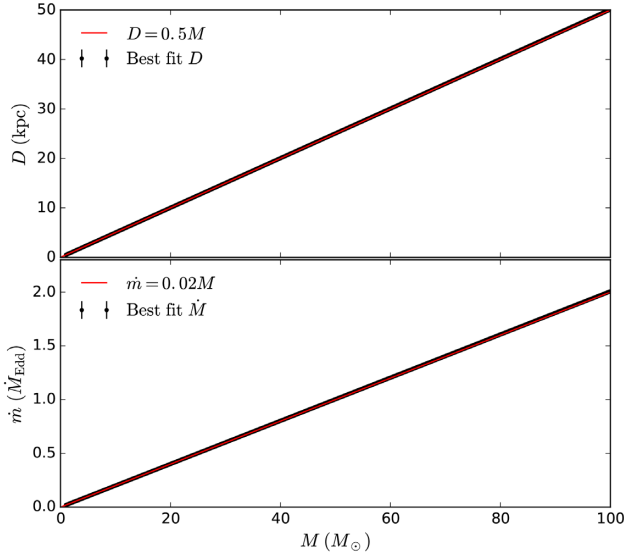


Figure 8. Best-fitting distances and accretion rates as a function of mass for a single *NuSTAR*–XRT spectrum. Both are completely degenerate with M , but the ratio between the parameters is fixed and very strongly constrained.

Table 2. The effect of fitting a well-constrained spectrum with different colour-correction factors. Units are as in Table 1. The right column gives the standard deviation of the three values.

Parameter	$f = 1.4$	$f = 1.7$	$f = 2.0$	σ
a	>0.999	>0.993	$0.81^{+0.03}_{-0.06}$	0.1
i	49.4 ± 0.1	38.7 ± 0.1	51^{+4}_{-2}	6
\dot{M}	0.237 ± 0.001	0.253 ± 0.002	$0.56^{+0.08}_{-0.04}$	0.18
χ^2	585	324	325	

4.1 Colour-correction factor

Arguably the largest source of systematic error in continuum fitting measurements is the colour-correction (or spectral hardening) factor, f (Shimura & Takahara 1995). This converts between the colour temperature and the effective temperature of the disc, and depends on the accretion rate and structure of the disc. The colour-correction factor, f , is set as 1.7 by default but could be as low as ~ 1.4 or as high as ~ 2 (Davis et al. 2005; Davis & Hubeny 2006; Reynolds & Miller 2013), and is dependent on the mass and accretion rate of the black hole (Davis & El-Abd 2018). We test the likely impact of assuming the wrong value of f by fitting the same well-constrained spectrum as before ($M = 10 M_\odot$, $D = 3$ kpc, $a = 0.998$, and $i = 40^\circ$), fixing $f = 1.4$ and 2.0 . This has a moderate impact on the spin and inclination, giving a systematic error of $\sim \pm 0.1$ in spin and $\sim \pm 10^\circ$ in inclination, and a larger effect on the accretion rate, although only in the high f case (the values are given in Table 2). The fit is also much worse in the low f case, but equivalent for high f . We also note that the colour-correction factor is meant to account for lower free-free absorption at high energies allowing a deeper view of the disc, but does not necessarily perfectly approximate this effect, which could have implications for precision spectroscopy.

4.2 Calibration uncertainties

Because the precision needed to make these measurements is extremely high, it approaches the limit of the *NuSTAR* instrumental calibration, which dominates over Poisson noise as the main source of uncertainty in the spectrum at high count rates. This is likely to be an issue for this technique with future instruments as well, although the exact nature will vary, and should in general be carefully considered. The calibration of *NuSTAR* is discussed extensively in Madsen et al. (2015), so here we focus on two main effects that are particularly relevant to this method:

(i) Uncertainty in column density to the Crab nebula. The nominal column density used in the *NuSTAR* calibration is $2 \times 10^{21} \text{ cm}^{-2}$, but it could be as low as $1 \times 10^{21} \text{ cm}^{-2}$ or as high as $6 \times 10^{21} \text{ cm}^{-2}$. To estimate the uncertainty introduced by this, we take a *NuSTAR* spectrum that correctly returns the spin and inclination values ($M = 10 M_\odot$, $D = 3$ kpc, $a = 0.998$, $i = 40^\circ$) and fix the column density at values $\pm 5 \times 10^{21} \text{ cm}^{-2}$ from the simulated value of $5 \times 10^{21} \text{ cm}^{-2}$. We find almost no effect on the fit parameters from this, with the largest effect being on inclination, where it adds a systematic error of $\pm 3^\circ$. We conclude that this is not a significant issue, although we note that the effect could be more important in sources where the intrinsic column is low and the uncertainty is thus fractionally larger.

(ii) Uncertainty in the location of the optical axis. The position of the optical axis is only known to within ~ 30 arcsec. The off-axis position is needed for accurate calculation of the effective area, which is crucial for measuring the flux. We test the effect of fitting with an auxiliary response file with the wrong off-axis position. We find that, even with the off-axis angle 1 arcmin larger than the true angle, the spin and inclination parameters are not significantly affected. The largest effect is at high energies, where the vignetting effects are strongest (Madsen et al. 2015, fig. 7) and result in broad line-like features at the ~ 1 – 2 per cent level in the spectrum. We conclude that this uncertainty in the optical axis is unlikely to have a significant impact on our results.

The other main source of uncertainty in the *NuSTAR* calibration are residuals due to spline fitting to the Crab spectrum. This typically produces narrow-line-like features of the order of ± 1 per cent, which should not have a major impact on the results of continuum fitting.

5 DISCUSSION

Arguably the most exciting potential use of being able to measure an additional parameter with continuum fitting is searching for warped accretion discs. A common prediction of accretion simulations is that the disc should bend or tear when the black hole spin axis and the binary axis are misaligned: the material accreted from the companion starts orbiting around the binary axis, but at small radii it aligns with the black hole (Bardeen & Petterson 1975). This is an important tracer of the black hole formation history, as the misalignment is thought to be produced by the supernova ‘kick’ as the black hole is born (see e.g. Jonker & Nelemans 2004, and references therein). Because X-ray measurements probe the inner accretion disc, a difference between the inclination measured using X-rays and the binary inclination may indicate a warp in the disc. However, to date all measurements of inner disc inclination with X-ray spectra have been made using only reflection spectroscopy. While there is generally good agreement between the spin values measured with this technique, in some sources there is significant variation in

inclination (see e.g. the list of inclinations measured for GX 339-4 listed in the introduction of P16). Inclination measurements with continuum fitting offer an independent method for establishing the inner disc inclination, which shares few systematic uncertainties or modelling assumptions with reflection spectroscopy. If both reflection and continuum methods indicate a different inclination from the binary, we can therefore be reasonably confident in the detection of a warped disc.

5.1 Additional systematic effects

The main assumption in both continuum and reflection methods is that the disc truncates at the ISCO. If the disc is either truncated or produces significant emission from the plunge region then the estimate of spin from both methods will be wrong. Truncation is not a significant worry in bright states with strong disc emission, and significant truncation can generally be ruled out when the measured spin is high. Zhu et al. (2012) considered the impact of emission from inside the ISCO on continuum fitting results. They find that the main effect of emission from the plunge region is to modestly raise the measured spin parameter. More worryingly, from our perspective, it also produces a weak power-law tail at high energies. Depending on the photon index of this tail (which is not well established), this could potentially be degenerate with the relativistic distortion of the disc spectrum, compromising the simultaneous spin/inclination measurements. A better understanding of this scattered tail is needed to rule this out. From the reflection perspective, simulations suggest that emission from the plunge region should be small (Reynolds & Fabian 2008), and contribute less to the total emission the higher the spin is, due to the smaller size of the region. This means that, while still potentially important, the effect of truncation or emission from the plunge region should be small when the measured spin is high.

We have assumed a simple disc plus power-law model for our spectra. As these measurements rely on high-precision measurements of spectral shape, they are highly sensitive to the exact continuum used. If additional spectral components or spectral complexity (such as the stratified Comptonization discussed in Kawano et al. 2017) are present in the soft state then these could be highly degenerate with the relativistic distortion of the spectrum. Similarly, overlap with the red wing iron line from relativistic reflection could also produce enough flux on the blue side of the disc spectrum to make such measurements impossible for certain spectra.

For our simulations we have used a modified version of the KERRBB model, which has been at least partly superseded by KERRBB2 and BHSPec. We chose to use KERRBB for the sake of simplicity and speed of calculation, as we use a large number of simulated spectra, and because we have only implemented higher order interpolation in this model (see Appendix A). We also simulate a small number of spectra with the other two models and confirm that the basic principle holds: with high enough signal, both a and i can be constrained.

A flaw in this method as used in P16 was the modelling of the Compton scattered tail. We used the COMPTT (Titarchuk 1994) model for the continuum, which does not subtract the photons scattered into the tail from the disc flux, so the photon flux is not conserved. As the continuum-fitting method relies in part upon measuring the total flux, this can potentially lead to an underestimation of the spin (or, equivalently, an overestimation of inclination). To avoid this problem, Steiner et al. (2009a) created the *simpl* model, which we use for the majority of this work. The model

simpl self-consistently removes photon flux from the disc spectrum and adds it to a power-law tail, which is cut off at low energies (standard power-law models do not include this, overpredicting the flux at energies below the peak of the disc spectrum). The disadvantage of this model is that it struggles to reproduce more complex continua, which is much more of an issue with *NuSTAR* than with softer or less sensitive detectors. In P16 we found that *simpl* was not able to reproduce the curvature of the Comptonized component, and relied on COMPTT instead. Now, however, Steiner et al. (2017) have released the SIMPLCUT model, which allows the high-energy cut off to be modelled either with an exponential function or with a physical Comptonization turnover. This model removes the conflict between modelling the high-energy spectrum accurately and conserving photon flux.

5.2 Mass and distance measurements

We showed in Section 3.3 that M , D , and \dot{m} are completely degenerate, but also that their ratios are fixed and strongly constrained. This is contrary to the result of P16, where we claimed a measurement of $M = 9.0^{+1.6}_{-1.2} M_{\odot}$ and $D = 8.4 \pm 0.9$ kpc. The most plausible reason for this disagreement is that we underestimated the error in these parameters in P16, due to not fully exploring the degeneracy between them in our MCMC analysis. This is due to the narrowness of the degeneracy (i.e. the huge increase in χ^2 away from the constant relation), which means MCMC steps will only be accepted if they are almost directly along the line of the degeneracy. This results in a very slow exploration of the parameter space, leading us to conclude that the chains had converged when they had not. However, as the ratios between the three parameters are fixed, we can revise that result to conclude that the true mass and distance should lie close to the line $D/\text{kpc} = 0.93M/M_{\odot}$. We have since continued the MCMC run for 10^6 additional steps, and find that it performs as expected: very slowly expanding along the narrow degeneracy. The other parameters are not affected by this, having fully converged in the original run, and are entirely consistent with the values reported in P16.

While we now know that simultaneous constraints on M , D , and \dot{m} from X-ray spectroscopy alone are not possible, this does not mean that there is no prospect for measuring these values. We need only constrain one parameter through other means to establish the other two to a high degree of precision. Trigonometric parallax (for example with radio telescopes or with *Gaia*) can be used to measure precise distances, without corresponding mass measurements. These distances could then be combined with X-ray spectroscopy to accurately measure the mass. Similarly, with future instruments it will be possible to perform detailed X-ray spectroscopy of extragalactic XRBs, where the distance is known to a high degree of relative precision but the mass cannot be measured dynamically. For these sources, X-ray mass measurements could be a very powerful tool for studying their mass distribution.

Alternatively, we can simultaneously measure M and D if the accretion rate is known. One potential method to establish \dot{m} is to base it on the state transitions transient XRBs go through, which are thought to occur when the accretion rate reaches a certain level (e.g. Maccarone 2003). Although the physics here is uncertain, it should be possible to give rough estimates for many transient sources, which can then be used to approximate M and D .

A final way of overcoming this difficulty is to use time-resolved spectroscopy. As the mass and distance do not change significantly with time, any two spectra (of the same state) with different luminosities must have different accretion rates. By fitting two

or more spectra from the same source simultaneously, it may be possible to constrain the difference in m , and hence infer M and D .

6 CONCLUSIONS

(i) *NuSTAR* spectra can simultaneously constrain both spin and inclination, provided that the source is bright enough and enough of the disc photons are in the *NuSTAR* band. This conclusion applies in principle to all instruments with sufficient sensitivity and resolution to measure the relativistic distortion of the disc spectrum, which is strongest above 5 keV.

(ii) These constraints are only strong for bright, high spin sources. For low spins and fluxes, the degeneracy between a and i is unbroken, so the spin is not strongly constrained. This error is well described by the measurement errors, so these measurements should be reliable.

(iii) The addition of low-signal soft X-ray (e.g. *Swift* XRT) data to *NuSTAR* significantly improves the accuracy of the result. While the XRT does not contribute to measuring the relativistic distortion, it does give an accurate measurement of the peak of the blackbody flux, which is outside the *NuSTAR* band in many cases.

(iv) The mass and distance parameters cannot be simultaneously measured while the accretion rate is free. Fixing a and i to values taken from e.g. reflection spectroscopy and fitting for M and D gives a perfect degeneracy between the two, where the ratio is constant. This degeneracy could be broken by fixing the distance to that measured using *Gaia* or inferring an accretion rate from state transitions. Further into the future, this could be used to measure masses in extragalactic sources, where dynamical masses cannot be measured but the distance is known precisely.

(v) We consider the impact of various systematic effects on simultaneous spin and inclination constraints from continuum fitting. We find that the instrumental calibration uncertainties associated with *NuSTAR* have minimal impact on the results, but that uncertainties in spectral modelling, in particular complex spectral curvature or additional continuum components, can have a major impact on the measured parameters.

Our results show that there is no reason why spin and inclination cannot in principle be measured simultaneously using the disc spectrum in XRBs. There are potentially large systematic errors in the results caused by biases in the spectral fitting, but crucially these are largely independent of the systematic errors in reflection fitting (and also are not unique to this usage of the continuum method). This means that independent measurements of the spin and inclination can be taken with the two methods and used to reliably estimate the inclination of the inner accretion disc, providing a reliable means to search for disc warps.

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Zhang S. N., Cui W., Chen W., 1997, *ApJ*, 482, L155Zhu Y., Davis S. W., Narayan R., Kulkarni A. K., Penna R. F., McClintock J. E., 2012, *MNRAS*, 424, 2504**APPENDIX A: POLYKERRBB**

Both KERRBB and BHSPec are calculated on a grid of parameter values, which are then interpolated between. For KERRBB the interpolation is done by the model code, for BHSPec the table is interpolated by XSPEC, in both cases linear interpolation is used. As we have to make high-precision measurements of the spectral shape in order to determine both the spin and inclination, the limited resolution of these grids becomes an important factor.

In Fig. A1 we show the χ^2 spin contour for POLYKERRBB with different degrees of interpolation, for one of the simulated spectra ($a = 0.998$, $i = 40^\circ$, $M = 10 M_\odot$, $D = 3 \text{ kpc}$). The grid points where KERRBB is evaluated are marked by the vertical lines, and between the grid points are a series of false minima or maxima, depending on the degree of interpolation. While the best fit is close to the true value in all cases, the fit could easily become trapped in one of these minima as they are both deep and wide compared to the typical step size. The effect of this is twofold: first, the fit could get stuck in a false minimum and return the incorrect value, and secondly the errors measured from such a contour are likely to be significantly under- or overestimated.⁴ Adding a systematic error to the data points scales the whole contour, and does not remove the false minima/maxima.

⁴This has much less effect on results calculated with a fixed inclination, where the constraint on spin is orders of magnitude stronger. Adding the contour with a fixed i value to Fig. A1 gives a line that is almost vertical at the correct value.

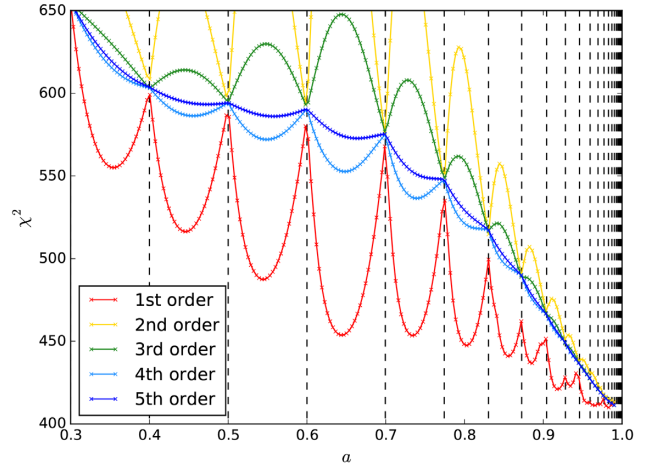


Figure A1. χ^2 contour for simulated POLYKERRBB spectra ($a = 0.998$, $i = 40^\circ$, $M = 10 M_\odot$, $D = 3 \text{ kpc}$). Crosses indicate where the fit is evaluated, and vertical dashed lines show the values of spin used in KERRBB. The lowest points (red) are for the 1st order interpolation, and correspond to standard KERRBB.

With a fifth-order polynomial interpolation the false minima are largely removed, although some residuals are still present and do slightly bias our best-fitting results towards being found between bins. For our purposes, this model is sufficient, however a more sophisticated spline interpretation would presumably offer an improvement. We note that issues with sparse grid points and linear interpolation in table models are not unique to KERRBB (indeed, XSPEC exclusively uses linear interpolation for table models), and are likely to be of increasing importance as the sensitivity of X-ray spectra increases.

The modified POLYKERRBB code is available from github.com/dbuisson/polykerrbb.

This paper has been typeset from a \LaTeX file prepared by the author.